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The pair is always difficult to measure, even with the 36-inch refractor, for the maximum distance between the components is only $0''.17$, the minimum value being $0''.07$. For this reason, probably, measures have not been made, so far as I am aware, by any other observers, and my orbit necessarily rests entirely upon my own observations.

The elements given are to be regarded as preliminary only, but the general form of the apparent ellipse is well defined. The orbit represents my measures as well as could be expected in a pair of this character, the residuals in distance falling within the limits $\pm 0''.03$, and the residuals in angle, reduced to arc measure, being considerably smaller (except for the year 1904, for which the residual is $0''.05$).

Only three other visual binary stars— δ *Equulei*, 13 *Ceti* and κ *Pegasi*—are known to have shorter periods of revolution than that found for A 88.

R. G. AITKEN.

December, 1912.

POSITION OF NOVA GEMINORUM, 1912.

Two observations of last year's *Nova* made with the meridian circle on January 27th and 29th, give the position referred to mean equinox 1913.0:—

R. A. $6^h 49^m 15^s.773$; Decl. $+ 32^\circ 15' 1''.64$; Epoch 1913.08.

Two observations were obtained last year, immediately following the receipt of the announcement of discovery in March. Referred to mean equinox 1912.0, the position was:—

	$6^h 49^m 11^s.796$; $+ 32^\circ 15' 5''.63$; Epoch 1912.22	
Precession	$+ 3.896$	$- 4.27$
1913.0	15.692	1.36

Comparison of the results, differing ten months in epoch, gives no indication of change of position. The drop in brightness, from fifth magnitude in March, 1912, to eighth magnitude now, would produce an effect upon the observed transit of the *Nova* of about one third of the right ascension difference, due to magnitude equation (three units at $0^s.008$), and the outstanding differences are not greater than might be due to probable errors of observation.

R. H. TUCKER.